

## Gamma ray signals from cosmic ray scattering on axionlike particles

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Dark matter (DM) may be comprised of axionlike particles (ALPs) with couplings to photons and the standard model fermions. In this paper, we study photon signals arising from cosmic ray (CR) electron scattering on background ALPs. For a range of masses we find that these bounds can place competitive new constraints on the ALP-electron coupling, although in many models lifetime constraints may supersede these bounds. In addition to current Fermi constraints, we also consider future e-Astrogram bounds which will have greater sensitivity to ALP-CR induced gamma-rays.

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### I. INTRODUCTION

The nature of the nonluminous dark matter (DM) is one of the greatest mysteries in physics. At present, everything that is known about DM has been derived on the basis of its gravitational interactions with ordinary matter. The race is on to detect any nongravitational interactions of DM, which would be a significant step forward in our understanding of the most abundant type of matter in the Universe.

A potential particle candidate for the DM is the axion, a light pseudoscalar which emerged as a means of solving the strong  $CP$  problem of QCD [1–3], but can also provide the necessary DM component, thus neatly solving two outstanding problems simultaneously (for a recent review of QCD axions, see [4]). Outside of attempts motivated by a solution the strong  $CP$  problem of QCD, light pseudoscalar particles known as ALPs—axion-like particles—have undergone intense investigation recently as they naturally appear in theories of physics beyond the standard model, e.g., string theory or other beyond the standard model theories, as the pseudo Nambu-Goldstone boson resulting from the breaking of a  $U(1)$  symmetry. This can induce a photon coupling through a mixing term of the form

$$\mathcal{L} \supset -g_{a\gamma} a F_{\mu\nu} \tilde{F}_{\mu\nu}, \quad (1)$$

where  $F_{\mu\nu}$  is the electromagnetic field strength tensor with  $\tilde{F}_{\mu\nu}$  as its dual,  $a$  is the ALP field and  $g_{a\gamma}$  is a coupling constant with units of  $\text{mass}^{-1}$ . In addition, ALPs may also couple to electrons via

$$\mathcal{L} \supset -C_{ee} \frac{m_e}{\Lambda} a \bar{e} i \gamma_5 e, \quad (2)$$

with the dimensionless ALP electron coupling  $C_{ee}$ . We will study the implications of both couplings in this paper. While specific models of the QCD axion constrain the axion parameter space of allowed couplings and masses due to relations between the scale at which the new symmetry introduced to solve the strong  $CP$  problem is broken, the axion mass, and the couplings within specific models [5–8], ALPs relax these mass and coupling relations, allowing for a broader allowable parameter space.

A tremendous variety of experiments have been specifically designed to search for the axion-photon coupling including helioscopes such as CAST [9,10] and IAXO [11,12], haloscopes including Abracadabra [13,14], ADMX [15,16], CASPEr [17], HAYSTAC [18,19], light-shining-through-walls experiments including ALPSII [20], and interferometers [21,22] such as ADBC [23] and DANCE [24]. Beam dump experiments can also manifest sensitivity to axion-photon couplings through axion decays or bremsstrahlung, including FASER [25], LDMX [26,27], NA62 [28], SeaQuest [29], and SHiP [30]. Additionally,

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there are hybrid beam dump/haloscope designs such as PASSAT [31]. A summary of recent collider constraints can be found in [32]. Neutrino and dark matter experiments such as NOMAD [33], XMASS [34], EDELWEISS-III [35], LUX, [36], PandaX-II [37], Xenon1T [38,39], and SuperCDMS [40] have been leveraged as axion searches. Geoscopes have also been proposed to study axion-electron couplings [41]. The axion-photon coupling has been explored in direct detection experiments such as DAMA [42], EDELWEISS-II [43], XMASS [44], and Xenon1T [39]. Axion-nucleon couplings are investigated from a solar axion flux produced through nuclear transitions, and can also be searched for through resonant absorption by laboratory nuclei [45–50] including the CUORE Collaboration [51,52], and the proposed GANDHI experiment [53]. The reactor experiment TEXONO has produced bounds on axion couplings [54], and upcoming neutrino experiments at nuclear reactor sites also having projected sensitivity in interesting axion mass and coupling parameter regions [55,56]. ALPs are also investigated in various astrophysical objects such as galaxies, galaxy clusters, quasars, pulsars, extragalactic space etc. utilizing the Primakoff conversion and/or in the presence of the magnetic field associated with these objects, e.g., [57–66].

Searches for ALPs can be direct, taking place in laboratories on Earth as listed above, or indirect, looking for the impact of ALPs on astrophysical objects. Traditional indirect detection searches of dark matter focus on its decay and/or annihilation products. However, the scattering of dark matter on electrons or nuclei is another possible probe [67].

In this paper we perform an indirect search by using measurements from the Fermi Gamma-Ray Space Telescope [68] and the Energetic Gamma Ray Experiment Telescope (EGRET) to constrain possible ALP interactions. We examine both electron and photon couplings to ALPs and find that the electron couplings produce bounds which are potentially competitive with bounds from other experiments. Additionally, we calculate future constraints from the planned E-Astrogram gamma-ray observatory.

## II. CALCULATIONAL FRAMEWORK

In this paper we focus on axionlike particles (ALPs) as dark matter and examine the utility of cosmic ray-ALP scattering to probe the coupling of ALPs to ordinary matter.

First notice that for a given incoming cosmic-ray energy, there is a minimum and maximum gamma-ray energy produced

$$E_\gamma^{\max/\min} = \frac{m_a^2 + 2m_a E_e}{2E_e + 2m_a \mp 2\sqrt{E_e^2 - m_e^2}}, \quad (3)$$

where  $m_a$  is the ALP mass,  $E_e$  is the electron energy, and the maximum (minimum) occurs for forward (backward) scattering. We show in Fig. 1 the dependence of the CR electron energy on the produced gamma-ray energy. Notice

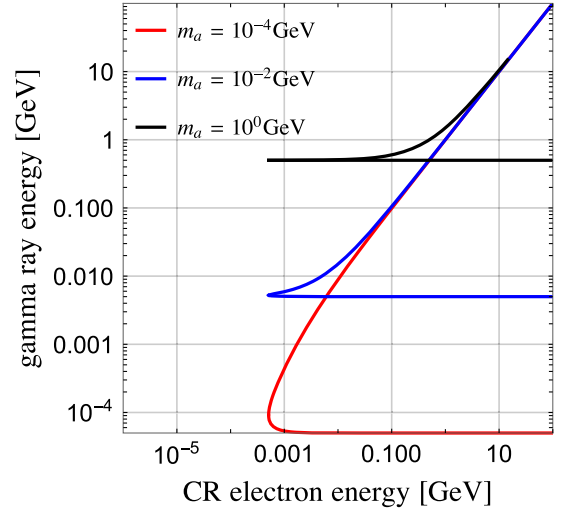


FIG. 1. Here we plot Eq. (3) in order to illustrate how the outgoing photon energy is influenced by the incoming CR energy and the ALP mass. At each CR energy the outgoing photon energy is constrained to lie to the right of the contour for fixed ALP mass.

that at the highest energies, the gamma-ray energy is simply equal to the incoming CR electron energy. However for a fixed ALP mass, the gamma-ray energy will always be larger than half the ALP mass,  $E_\gamma \geq m_a/2$ .

Having addressed the kinematics of ALP induced gamma-rays, we move on to the dynamics. We compute the expected gamma-ray flux from ALP-electron scattering as

$$\frac{d\Phi_\gamma}{dE_\gamma} = \int_V dV \int_{E_e^{\min}} dE_e \frac{d^2\Gamma_{e\rightarrow\gamma}}{dE_e dE_\gamma} \quad (4)$$

$$= D_{\text{eff}} \frac{\rho_{\text{DM}}}{m_a} \int_{E_e^{\min}(E_\gamma)} dE_e \frac{d\sigma_{e\gamma}}{dE_\gamma} \frac{d\phi_e^{\text{LIS}}}{dE_e}, \quad (5)$$

where  $\rho_{\text{DM}} \simeq 0.3 \text{ GeV cm}^{-3}$  is the DM density,  $\frac{d\phi_e^{\text{LIS}}}{dE_e}$  is the electron flux of the local interstellar spectrum (LIS), and  $D_{\text{eff}}$  is the effective distance over which we include CRs as the source of ALP induced gamma-rays. For a sphere of radius 1 kpc centered on the Earth, and assuming an NFW profile, we take  $D_{\text{eff}} \simeq 1 \text{ kpc}$  [69]. Note that the lower bound on the integral,  $E_e^{\min}(E_\gamma)$  is the minimum electron energy to produce a given gamma-ray energy. This quantity can be found by inverting Eq. (3). We use the LIS fluxes from Ref. [70,71], which are obtained from a fit to Voyager 1 [72], AMS-02 [73], and PAMELA data [74].

We will assume that only either the electron or the photon coupling to ALPs is present at low-energies. This is a conventional approach adopted in phenomenological studies of ALPS (e.g., [75]). Further, notice that although the electron coupling will induce the photon coupling at

one-loop, this coupling will be suppressed by a factor  $\mathcal{O}(\alpha_{\text{EM}})$ . As a result, we will ignore this contribution since it is not expected to change our results appreciably. Note that it is possible to construct viable scenarios where the ALP-photon coupling can be minimized, see for example [76–78].

We will examine the scattering process  $a(k) + e(p) \rightarrow \gamma(q) + e(p')$  with only the pseudoscalar  $g_{ae}$  coupling being nonzero. Performing the traces and computing all the scalar products, the squared amplitude for ALP-electron scattering is

$$\frac{1}{2} \sum_{\text{spins}} |\mathcal{M}|^2 = g_{ea}^2 e^2 \left[ -\frac{2(m^4 - m^2(2m_a^2 + s + u) + su)}{(u - m^2)^2} - \frac{2(m^4 - m^2(2m_a^2 + s + u) + su)}{(s - m^2)^2} + \frac{4(m^4 + 3m_a^2 m^2 - m^2(2s + t) + (s - m_a^2)(s + t))}{(s - m^2)(u - m^2)} \right], \quad (6)$$

where the 1/2 factor arises from averaging over the initial electron spin (the axion has spin zero, thus only contributes a factor of one to the counting), and we have used the Mandelstam variable relation  $s + t + u = 2m^2 + m_a^2$  to simplify the result.

Next, in the case of nonzero photon-ALP coupling we find that the matrix element for  $e + a \rightarrow e + \gamma$  is

$$\frac{1}{2} \sum_{\text{spins}} |\mathcal{M}|^2 = \frac{e^2 g_{a\gamma\gamma}^2}{t^2} (m_a^2 t(m^2 + s) - m_a^4 m^2 - t((s - m^2)^2 + st) - t(t - m_a^2)^2/2). \quad (7)$$

This matches with the results of Appendix B of [79] when using the Mandelstam variable formulation.

With the matrix element squared in hand, in standard fashion, we then obtain the cross section for each process via

$$\frac{d\sigma}{dE_\gamma} = \frac{1}{32\pi m_a |\vec{p}|^2} \left( \frac{1}{2} \sum_{\text{spins}} |\mathcal{M}|^2 \right), \quad (8)$$

where we have included the 1/2 factors as in Eqs. (6) and (7). With these results in hand, we will turn to the observational constraints on these scattering processes.

### III. RESULTS

We now confront the above predictions for ALP produced gamma-rays with experimental observations. Given that we need both large DM and CR densities, the galactic center region would be an obvious target. However, the strength of the ALP-CR interactions we are interested in are particularly large, and these may alter the ordinary expectation of cuspy density profiles. In addition, the signal-to-background ratio is presumably worse in this region given the large number of SM processes leading to photon production in this high-density region.

For the reasons argued above, we follow Ref. [67] and utilize high-latitude data ( $|b| > 20^\circ$ ) to constrain background ALPs. In this region the differences between cuspy

and cored DM density profiles are less relevant, and the contributions from ordinary SM processes are suppressed. Observations of the isotropic gamma-ray background (IGRB) are a useful dataset in this regard. The IGRB is the residual component of the extragalactic background with diffuse galactic contributions and known point sources removed. This flux has been measured by the Fermi collaboration [68] between (0.1 and 820) GeV. From statistical uncertainty alone, the total intensity attributed to the IGRB was found to be  $(7.2 \pm 0.6) \times 10^{-6} \text{ cm}^{-2} \text{ s}^{-1} \text{ sr}^{-1}$ . However, the collaboration estimates an additional (15–30)% systematic component driven by uncertainty in Galactic diffuse foregrounds. Lastly, we note that Ref. [80] has also used IGRB data to constrain DM, albeit in an annihilating DM framework.

To be conservative, we report bounds from Fermi by simply requiring that the predicted ALP induced gamma-ray flux not exceed the flux measured by Fermi. This is even more conservative than the approach adopted in Ref. [67]. There it was argued that the new physics flux could at most comprise 20% of the IGRB since roughly 80% can be explained with known astrophysics involving nonblazar active galaxies [81] and star-forming galaxies [82].

Our results from the Fermi data [68] are shown in Fig. 2 and Fig. 3. As can be seen, a region of  $g_{ae}$  parameter space at relatively large ALP mass can be excluded which is not covered by existing experimental or observational data. It should be noted, however, that this region has sufficiently large couplings that would naively lead to very rapidly decaying ALP DM, via  $a \rightarrow \gamma\gamma$ . Although the goal of the present work is not to construct a complete ALP model, it should be noted that various “photophobic” scenarios [76,83] have been discussed in the literature which may alleviate such short lifetimes. We note that the bounds at low ALP mass can be improved by extrapolating the CR flux to higher energies. Thus with future high-energy CR data we expect the bounds at low ALP mass to improve considerably.

Next one may wonder how much these constraints can be improved with future data. As a proof of principle, we look

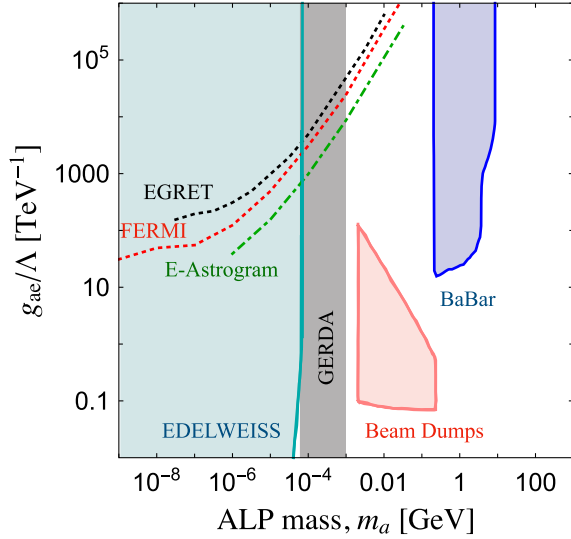


FIG. 2. Here we show our conservative constraints from Fermi data [68] on the ALP-electron coupling, along with a variety of additional constraints from Babar [75,84], EDELWEISS-III [35], and GERDA [85], and beam dumps [86]. Note that while the EDELWEISS and GERDA bounds assume that ALPs comprise the DM, the beam dump and *BABAR* bounds are independent of this assumption.

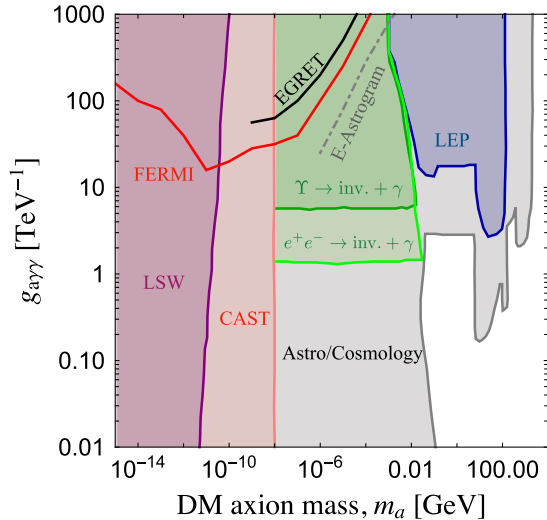


FIG. 3. Here we show the bounds on the axion-photon coupling from Fermi and EGRET, along with a projected sensitivity for E-Astrogram.

at the expected sensitivity of E-Astrogram [87]. E-Astrogram will dramatically improve on existing measurements in terms of statistics, but should still measure approximately the same central values of the photon flux in each bin already detected by other experiments. Thus in

E-Astrogram’s energy range 0.3 MeV to 3 GeV, we assume that the measured photon flux will have the same central values as already measured. Further, since this is an optimistic projection we will also assume that ordinary astrophysical explanations of known physics can be found for 90% of the measured flux, such that the ALP induced flux can at most account for 10% of the future IGRB flux seen by E-Astrogram. These projected constraints are shown in Fig. 2 and Fig. 3.

Lastly while we have focused on IGRB data to constrain ALP-CR interactions, one could potentially use other datasets. If for example a whole-sky search was conducted, this could enhance the signal by roughly a factor of  $\sim 5$ . Optimistically then, such a search may be able to improve over the bounds we have found here by a factor of  $\sim 1/\sqrt{5} \simeq 0.45$ .

#### IV. CONCLUSIONS

We have investigated for the first time the production of photons from ALP-cosmic ray collisions. Through the use of FERMI gamma-ray observations of the isotropic gamma-ray background, we have found this process provides new competitive constraints on ALP-electron couplings, including probing new parameter space in the  $g_{ae} - m_a$  plane. While here we have focused on the electron and photon couplings, cosmic rays could also be used to develop new constraints on the ALP-proton couplings.

In addition to examining the bounds implied by present FERMI data on the IGRB, we have also made optimistic projections for the sensitivity of the future E-Astrogram observatory to ALP-induced gamma-rays. We have been conservative here in that we systematically underestimate the photon flux since we have only included diagrams with single outgoing photons. Second, we have only utilized existing data from Fermi, though one could exploit the predicted spatial morphology of the CR-ALP induced flux to impose some spatial cuts which could improve the signal-to-background ratio. Lastly we note that with future improvements in our understanding of the IGRB these bounds may be considerably improved.

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